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Odin/OSIRIS observations of stratospheric BrO: Retrieval methodology, climatology, and inferred Br$_y$

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A 7+ year (2001–2008) data set of stratospheric BrO profiles measured by the Optical Spectrograph and Infra-Red Imager System (OSIRIS) instrument, a UV-visible spectrometer measuring limb-scattered sunlight from the Odin satellite, is presented. Zonal mean radiance spectra are computed for each day and inverted to yield effective daily zonal mean BrO profiles from 16 to 36 km. A detailed description of the retrieval methodology and error analysis is presented. Single-profile precision and effective resolution are found to be about 30% and 3–5 km, respectively, throughout much of the retrieval range. Individual profile and monthly mean comparisons with ground-based, balloon, and satellite instruments are found to agree to about 30%. A BrO climatology is presented, and its morphology and correlation with NO$_2$ is consistent with our current understanding of bromine chemistry. Monthly mean Br$_y$ maps are derived. Two methods of calculating total Br$_y$ in the stratosphere are used and suggest (21.0 ± 5.0) pptv with a contribution from very short lived substances of (5.0 ± 5.0) pptv, consistent with other recent estimates.


1. Introduction

The importance of the inorganic bromine family (Br$_y$ = Br + BrO + HBr + HOBr + BrONO$_2$ + BrCl) has been recognized since the 1970s [Spencer and Rowland, 1978] and further elucidated in the 1980s [McElroy et al., 1986]. In the lower stratosphere, catalytic cycles involving bromine are important contributors to ozone destruction, particularly in the midlatitudes. By way of example, one such cycle is

(R1) BrO + O → Br + O$_2$

(R2) Br + O$_3$ → BrO + O$_2$

which results in a net loss of two odd-oxygen (O+O$_3$) molecules. Other cycles involving HOBr and ClO are well-documented [e.g., Chartrand and McConnell, 2000] and are not provided here.

Stratospheric Br$_y$ originates from the cross-tropopause transport of bromine-bearing species emitted from the surface. These sources can roughly be classified as (1) methyl bromide (CH$_3$Br), the single largest source with both natural and anthropogenic components; (2) halons (CBrClF$_2$, CBrF$_3$, CBr$_2$F$_2$, BrCBrF$_2$, and CBr$_2$F$_2$), strictly an anthropogenic source; and (3) bromine-bearing very short lived substances (VSLS) such as CH$_3$Br$_2$ and CHBr$_3$, with many other possible contributors [World Meteorological Organization (WMO), 2007]. It is this last source that is the subject of much debate as current estimates of their contribution to the stratospheric budget vary from 1.5 to 8 pptv [Pfeilsticker et al., 2000; Sinnhuber et al., 2005; Sioris et al., 2006; Kovalenko et al., 2007; Laube et al., 2008; Dorf et al., 2008; Hendrick et al., 2008]. The difference between the lower and upper estimate has a large impact on the importance of Br$_y$ on midlatitude ozone depletion [Salawitch et al., 2005].

While profiles of stratospheric HBr were measured from balloon-borne instruments in the 1990s [Johnson et al., 1995; Nolt et al., 1997], BrO is the only member of the Br$_y$ family that has been systematically and globally measured in the stratosphere, a fact that has complicated the closure of the Br$_y$ budget. A very recent exception to this are profiles of BrONO$_2$ observed by the Michelson Interferometer for Passive Atmospheric Sounding (MIPAS) satellite instrument [Höpfner et al., 2009]. In situ measurements of BrO exist...
[Brune et al., 1988], although these are sparse. Global measurements of the total vertical column density (VCD) of BrO have been made since 1995 by a series of nadir-viewing satellite instruments beginning with Global Ozone Monitoring Experiment (GOME) [Richter et al., 1998]. Ground-based zenith-sky observations of scattered sunlight have been made since the 1980s [Solomon et al., 1989] but, like nadir measurements, yield limited information on the vertical distribution of BrO. Stratospheric BrO profiles from balloon-borne instruments such as Systeme d’Analyse par Observations Zenithales (SAOZ)–BrO [Pundt et al., 2002] and Limb Profile Monitor of the Atmosphere/Differential Optical Absorption Spectroscopy (LPMA/DOAS) [Dorf et al., 2006] have been made semiregularly since the 1990s and for much of the time since remained the only source of profile information. More recently, satellite instruments such as Scanning Imaging Absorption Spectrometer for Atmospheric Chartography (SCIAMACHY) on Envisat [Bovensmann et al., 1999], in limb-scattering mode [Sinnhuber et al., 2005; Sioris et al., 2006; Hendrick et al., 2009], and Microwave Limb Sounder (MLS) on Aura [Kovaleenko et al., 2007], have been measuring BrO profiles. However, despite this additional information, the uncertainties in the VSLS contribution to stratospheric Br remain. The analysis of the same data are even leading to significant, and as yet unresolved, differences in the VSLS estimate [WMO, 2007; A. Rozanov et al., Retrieval of BrO vertical distributions from SCIAMACHY limb measurements, submitted to Atmospheric Measurement Techniques and Discussions, 2010].

In this work a new source of stratospheric BrO measurements is introduced, from the Optical Spectrograph and Infra-Red Imager System (OSIRIS) satellite instrument [Llewellyn et al., 2004], which should aid in our understanding and quantification of this important trace gas.

2. OSIRIS on Odin

Odin was launched in February 2001 into a 600 km circular, Sun-synchronous, near terminator orbit at an inclination of 97.8° and an ascending node at 1800 local solar time (LST) [Murtagh et al., 2002]. From launch until May 2007, Odin was a combination astronomy-aeronomy mission with time divided equally between the two modes. Since then Odin has operated exclusively in aeronomy mode. Odin carries two instruments: the Submillimetre and Millimetre Radiometer (SMR) [Frisk et al., 2003] which measures profiles of N\textsubscript{2}O, HNO\textsubscript{3}, O\textsubscript{3}, and ClO [Urban et al., 2005] and OSIRIS, the focus of this study. The instruments are coaligned and scan the limb of the atmosphere over a tangent height range of 7–70 km in approximately 85 s in a sawtooth pattern during normal stratospheric operations through controlled nodding of the satellite. Over an orbit, Odin makes roughly 65 such limb scans with roughly half in sunlight. Because of Odin’s orbit, OSIRIS does not provide profiles in the winter hemisphere, as solar zenith angles exceed 90°.

OSIRIS contains two optically independent components, the Optical Spectrograph (OS) and the Infra-Red Imager (IRI). The IRI is a three-channel camera, imaging the atmospheric airglow emissions near 1.27 μm and 1.53 μm in a limb–viewing tomographic mode [Degenstein et al., 2003]. The OS is a grating spectrometer that measures sunlight scattered from the Earth’s limb back into space. The limb-viewing geometry is illustrated in Figure 1. Specifically, OS measures limb-scattered radiance in the spectral range 280 nm to 800 nm at a spectral resolution of ∼1 nm (FWHM) at tangent height intervals of roughly 2 km. The detector is a CCD with 1353 × 32 pixels utilized. The width of a single pixel is about 0.4 nm in the spectral dimension, which gives a sampling ratio of 2.5. The instantaneous field of view (FOV) of the OS is 1 km in the vertical and 40 km in the horizontal at the tangent point. When the nodding of the spacecraft and the varying exposure time of the OS (∼0.01 s at 10 km increasing to ∼2 s at 50 km) are considered, the vertical resolution of the measured limb radiances is found to range from approximately 1 km at 10 km to 2 km at 50 km. For clarity, “OSIRIS” shall be used to reference the Optical Spectrograph.

Figure 1. Depiction of the limb-scattering geometry showing examples of single-scattered (SS), multiple-scattered (MS), and surface-scattered (AS) light paths. LOS denotes the OSIRIS line of sight, TH is the tangent height of the LOS, and TOA is top of atmosphere.
3. Retrieval Methodology and Characterization

This section describes the methods and algorithms used to create the OSIRIS BrO v4.0 data set.

3.1. Creation of Daily Zonal Mean (Level 1) Spectra

A preliminary evaluation of the OSIRIS spectra has revealed that there is generally an insufficient signal-to-noise ratio (SNR) for single-scan inversions of BrO by a factor of roughly 3. This has necessitated the use of co-added spectra, and sufficient SNR can be achieved with the averaging of 10+ spectra. To achieve this, the approach adopted was to create a daily zonal mean Level 1 (L1) product.

Limb scans from each day were collected into averaging bins: 10° wide in latitude and centered at 85°S to 85°N and sorted according to local time: AM (descending half of the orbit) or PM (ascending half of the orbit). The first step in the creation of the daily zonal mean L1 product was quality control (QC). A scan was discarded if (1) there were large (>6 km) gaps in tangent height, (2) it did not span 35 km, (3) Odin was in or near the South Atlantic Anomaly, and/or (4) the Moon was in the field of view. Individual tangent heights or scans were removed if a Level 1 flag or exception was triggered or there was uneven scanning (which could lead to additional noise and/or poor geolocation). Each scan that passed all QC tests was linearly interpolated onto a standard 81 point potential temperature grid spanning 275 to 5800 K, where the potential temperature was logarithm of this radiance ratio was performed between 38–46 km from the same zonal L1 scan. A spectral fit to the spectrum adopted is the mean over all tangent heights from 38–46 km. To achieve this, the approach adopted was to create a daily zonal mean Level 1 (L1) product.

3.2. Spectral Analysis

The BrO retrieval method is adapted from the OSIRIS NO2 algorithm [Haley et al., 2004; Haley and Brohede, 2007] and is similar to that used for the retrieval of OCIO from OSIRIS [Kreel et al., 2006]. Stratospheric profiles of BrO were retrieved in two steps, with the first being a spectral fit [Haley et al., 2004] to obtain so-called apparent slant column densities (SCDs), which represent the absorber-weighted path length through the atmosphere. The second step was an inversion of these SCDs into number density profiles using optimal estimation [Rodgers, 2000], described in section 3.5.

Tangent heights of 16–36 km were used in the inversion. Each daily zonal mean OSIRIS radiance spectrum in this altitude range is divided by a reference spectrum in order to remove Fraunhofer structure. The reference spectrum adopted is the mean over all tangent heights from 38–46 km from the same zonal L1 scan. A spectral fit to the logarithm of this radiance ratio was performed between 346–376.6 nm, which contains five BrO absorption features (the 5.0 to 1.0 bands) spanning 79 OSIRIS pixels. The basis functions used in the fit include BrO cross sections at 336.1, 376.6 nm, which contains five BrO absorption features (the 5.0 to 1.0 bands) spanning 79 OSIRIS pixels. The basis functions used in the fit include BrO cross sections at 336.1, 376.6, 403.9, 460.7, and 484.4 nm, and aerosol extinction, ozone at 203, 223, and 243 K [Bogumil et al., 2003], NO2 at 220 K [Volandel et al., 1998], and O3 (Belgian Institute for Space Aeronomy, http://www.aeronomie.be/spectrolab/O2.htm). In addition, a tilt/undersampling pseudo-absorber [Storrs et al., 2003; Haley et al., 2004] and a third-order closure polynomial were included. Sensitivity studies were conducted and found that including a correction for the Ring effect [Storrs et al., 2002] and polarization features [McLinden et al., 2002b] were unnecessary and so were excluded. All cross sections have been corrected for the so-called I0 effect [Allwell et al., 2002; Haley et al., 2004]. Also, a relative wavelength shift between the OSIRIS radiances and the basis functions is calculated and then applied [Haley et al., 2004].

An example of a representative spectral fit is shown in Figure 2. In this example, 20 individual scans were averaged over, and the root-mean-square (RMS) residual (over wavelength) is 6.7 × 10−4. To highlight the BrO absorption features, the fitted closure polynomial was removed from this spectrum. Upon averaging over a large number of residual (the difference between the fitted and measured) spectra, a common feature, or mode, emerged. The RMS of this common mode was about 3 × 10−4. To assess what impact it had on the fitted BrO amount, it was used as an additional basis function in the spectral fit.
found to be small, less than 3%, and appeared to be random. On this basis, it was concluded that inclusion of the mean residual in the spectral fit was not necessary, and not a significant source of error. An example of a SCD profile from 15 March 2003 at 55°N and its fitting uncertainty and RMS residual profile is shown in Figure 3.

3.3. Photochemical Box Model

[17] In this work the University of California, Irvine, photochemical box model [Prather, 1992; McLinden et al., 2000; Brohede et al., 2008] was employed to (1) calculate BrO a priori profiles required by the inversion (see section 3.5), (2) map the retrieved OSIRIS profiles from their local time to that of another measurement to facilitate a comparison, and (3) derive an estimate of total Br. In all applications, ozone, the NOy, Cly, and Bry families, and long-lived species (N2O, CH4, and H2O) were fixed. The remaining species were calculated to be in a 24 h steady state by integrating the model for 30 days with the diurnal cycle fixed on a specified Julian day. Photochemical rate data from the JPL 2006 compendium [Sander et al., 2006] were used.

[18] For the calculation of the a priori BrO, a set of tables were precomputed 3 times per month (for the 1st, 11th, and 21st day), every 2.5° in latitude, and at pressure altitudes between z* = 8 and 56 km in 2 km steps (where z* = −16log (p/1000) and p is pressure in hPa) with the entire diurnal cycle of BrO archived. The Bry abundance in the model was specified using the “organic” N2O-Bry correlation [Wamsley et al., 1998; Salawitch et al., 2005], adjusted to 2005 levels of N2O, with an additional 4 pptv to account for VSLS not included in the Wamsley et al. [1998] correlation, thus giving a total Bry of about 20 pptv. Monthly mean climatologies for temperature [Nagatani and Rosenfield, 1993] and ozone [McPeters et al., 2007] were used. The remaining species were either from three-dimensional model output (N2O, NOy) [Olsen et al., 2001] or tracer correlations (H2O,

Figure 2. Sample spectral fit for BrO from 15 March 2003, latitude of 55°N, AM, tangent height of 22 km (descending node; 0651 LST; SZA = 85.8°). (top) The fit to the measured OSIRIS differential optical depth (OD) spectrum. (bottom) The fitted BrO. The RMS residual (over wave length) is $6.7 \times 10^{-4}$, and the analyzed spectrum is the mean over 20 individual measurements. Each square denotes an OSIRIS pixel. To highlight the absorption features, the fitted closure polynomial has been subtracted.

Figure 3. Sample BrO slant column density (SCD) profile, the SCD uncertainty profile (standard error of the SCD) as determined from the spectral fit, and the RMS (over wave length) fitting residual from 15 March 2003, latitude of 55°N, AM (descending node; 0651 LST; SZA = 85.8°).
McLinden et al. [2000], four different sources in which partitioning, solved in a Gauss-Inversion at a SZA of 87° or smaller were considered.

The LIMBTRAN forward model, consisting of the model, the measured quantities, and some a priori information. This is an inverse problem, which uses the RT model, the measured quantities, and a priori information and the associated covariances. The a priori BrO was taken from the pre-calculated tables (section 3.3). An uncertainty of 100% was assigned to this. The off-diagonal elements in the a priori covariance matrix were constructed using a Gaussian correlation function with a constant correlation length (FWHM) of 4 km. This acts as a smoothing constraint to reduce measurement noise. It is noted that the 100% uncertainty and 4 km correlation length are ad hoc values and not in any way based on a measurement ensemble.

20] In the determination of BrO, the monthly mean fields were again used in the calculation of the BrO/BrO ratio required to convert BrO to BrO. Results from an investigation into the ability of the box model to simulate NOx partitioning through comparisons with the JPL MkIV interferometer measurements indicated good agreement [Brohede et al., 2008].

3.4. Radiative Transfer

21] A radiative transfer (RT) model capable of accurately simulating multiple scattering in limb geometry is required for any inversion of limb-scattered spectra. In this work the LIMBTRAN model [Griffioen and Oikarinen, 2000] was employed. For the simulation of a given OSIRIS scan, LIMBTRAN was initialized with a zonal mean temperature and neutral density profile from ECMWF analysis fields and ozone from a zonal mean climatology [McPeters et al., 2007]. In addition, a stratospheric aerosol extinction climatology [Bauman et al., 2003] was employed assuming a Heney-Greenstein phase function with an asymmetry parameter of 0.7, following Haley et al. [2004]. LIMBTRAN accounts for variations in SZA and the change in azimuthal angle (dAZ) along the tangent pathline of sight (LOS). When the change in SZA at the tangent point over the course of a scan becomes important, multiple LIMBTRAN runs were used to account for this. Simulated radiances were calculated on a 1 km vertical grid and at 1 nm increments (in contrast to NOx in which radiances were calculated every 2 nm). To reduce the CPU burden, weighting functions were calculated numerically at two wavelengths (355, 365.5 nm), in single-scattering, and assuming no aerosols.

22] While LIMBTRAN accounts for the variation of SZA along the LOS, it does not allow for any photochemical variability along the LOS. BrO experiences a large change through sunrise and sunset as a result of photochemistry and this may lead to a source of systematic error in the retrieval: the so-called diurnal effect [McLinden et al., 2006]. A research version of the retrieval algorithm using the Vector Orders-of-Scattering Radiative Transfer (VECTOR) model [McLinden et al., 2002a, 2006] does account for this, but the additional CPU requirements make this impractical for operational retrievals. As a result, only limb scans measured at a SZA of 87° or smaller were considered.

3.5. Inversion

23] The second step in the retrieval is the conversion of BrO SCDs derived from the spectral fit to number density profiles. This is an inverse problem, which uses the RT model, the measured quantities, and some a priori information. The LIMBTRAN forward model, consisting of the results of applying a spectral fit to RT model simulated radiances, was used to map the state-space (BrO profile) into the measurement space (the SCDs). In this study, the retrieval grid was chosen to be a 2 km altitude grid from 16 km to 36 km.

24] The inversion was accomplished using optimal estimation, or more specifically, the maximum a posteriori (MAP) estimator from Rodgers [2000], solved in a Gauss-Newton iterative way [Haley et al., 2004]. MAP is a Bayesian estimator giving the most probable solution based on the measurements and a priori information and the associated covariances. The a priori BrO was taken from the pre-calculated tables (section 3.3). An uncertainty of 100% was assigned to this. The off-diagonal elements in the a priori covariance matrix were constructed using a Gaussian correlation function with a constant correlation length (FWHM) of 4 km. This acts as a smoothing constraint to reduce measurement noise. It is noted that the 100% uncertainty and 4 km correlation length are ad hoc values and not in any way based on a measurement ensemble.

LIMBTRAN then was used to calculate synthetic limb radiances and weighting functions. The weighting functions, calculated using successive perturbations, describe how the forward model responds to a change in the BrO profile [Haley et al., 2004]. On the basis of these calculations, the observations, their uncertainties, and a priori information, MAP then determined an updated BrO profile which was used in LIMBTRAN for the next iteration. To avoid negative densities, a positive constraint was applied to the retrievals by inverting the natural logarithm of BrO number density. A discussion on how the transformation into logarithm-space affects retrievals and statistics is provided by Brohede et al. [2007b] in the context of OSIRIS NO2. In this case, no significant bias was found, and the shape of distribution was minimally affected. Similar conclusions are expected for BrO. Results from a sample measurement from 15 March 2004 (AM) at 45°N are shown in Figure 4.

3.6. Retrieval Characterization and Error Analysis

25] The MAP approach provides diagnostics that are very useful in characterizing the inverted profiles, including the averaging kernels, measurement response, and resolution. Examples of these are shown in Figure 4. The resolution, or the effective resolution of the inverted profile, was quantified as the full width at half maximum (FWHM) of the averaging kernels. The measurement response is given by the area of the averaging kernels and indicates the relative contribution of the measurements and a priori information to the retrieved profile. A value of 1 implies that the retrieved profile is based solely on the measurements, whereas a value of 0 indicates that the measurements had no contribution to the retrieved profile. Another important diagnostic is the information content of the measurements, which can be quantified by the degrees of freedom for signal, or DOFS [Rodgers, 2000], by calculating the trace of the averaging kernel. The DOFS may be interpreted as the number of independent pieces of profile information. This example has a DOFS of 5, and values are generally in the range of 4–6. Characteristics of the data set are shown in Figure 5.

27] As described by Rodgers [2000], four different sources of error can be identified: (1) smoothing error, (2) retrieval...
noise, (3) forward model error, and (4) forward model parameter error.

[28] The forward model error arises from an imperfect description of the physics in the RT model as well as approximations made in order to reduce the CPU requirements. Forward model parameter error is due to an imperfect knowledge of the parameters that impact the modeled limb radiances, including the aerosol, temperature, ozone, and neutral density profiles, surface albedo, and absorption cross sections.

[29] Quantifying the forward model and forward model parameter errors requires an off-line sensitivity analysis using noise-free synthetic measurements. For this a midlatitude atmosphere with a SZA of 85°, a change in azimuthal angle of 90°, and a surface albedo of 0.3 were specified. This follows the procedure used by Haley et al. [2004]. The forward model error was assessed by simulating spectra using more stringent model settings and then performing a retrieval using the standard settings. Aspects examined were vertical resolution, the use of convolved cross sections (as opposed to convolution of simulated high-resolution spectra), neglecting polarization, and the number of different SZAs simulated along the LOS. The total forward model error was determined by adding together the individual components. It

![Figure 4](image1.png)

**Figure 4.** BrO retrieval and diagnostics from 15 March 2003, latitude of 55°N, AM (descending node; 0651 LST; SZA = 85.8°): (a) number density (cm\(^{-3}\)); (b) volume mixing ratio (VMR); (c) relative retrieval error; (d) averaging kernels (AK) (solid) and retrieval response (dashed); and (e) retrieval resolution. The degrees of freedom for signal, a measure of the number of independent pieces of information, is 5 in this example.

![Figure 5](image2.png)

**Figure 5.** Characteristics of the OSIRIS BrO data set. Light grey lines are individual retrievals (every 100th profile shown), and thick black line is the median value considering all latitudes and both local times: (a) mixing ratio, (b) absolute retrieval error, (c) fractional retrieval error, (d) response, and (e) resolution.
is typically 3–5% and never exceeds 10% for BrO retrievals. The forward model parameter error was assessed by perturbing a model parameter by its typical uncertainty in the forward simulation and then performing a retrieval using the standard value. Again, the procedure of Haley et al. [2004] was followed. Here the individual sources were added in quadrature as they were assumed to be independent. Results are shown in Figure 6a. The largest sources are OSIRIS spectral resolution, or slit width, which was perturbed by 5% (this is discussed further in section 3.7) and BrO cross-section errors. In the error budget these two sources were treated separately from the other forward model parameter errors and thus were not included in the calculation of the total model parameter error curve in Figure 6a.

The forward model parameter error considered an error in the assumed ozone profile. Errors in the absorption spectra of interfering species, of which ozone is the most important, were not considered. This error source was difficult to assess as only an error in the spectral structure is relevant. Care was taken to minimize this error source by allowing for a relative shift in wavelength between the reference ozone spectra and OSIRIS. Further, ozone cross sections at three temperatures were included in the spectral fit. Preliminary sensitivity tests indicated that the use of an ozone cross section measured at a temperature that differs significantly from the effective atmospheric temperature may lead to profile errors of up to a factor of two.

The smoothing error arises because of the limited vertical resolution of the retrieval and the nonzero correlation length in the a priori covariance matrix. Retrieval noise represents the error source due to noise on the radiances, propagated through the retrieval, and is dependent on the a priori covariance matrix. The covariance matrix was constructed on the basis of box model simulations of BrO with an uncertainty of 100% and an ad hoc correlation length of 4 km, implemented as a smoothing constraint. As these were not based on an observed ensemble of states, the resultant smoothing error had little meaning. As a result we do not consider smoothing error in the error budget, but rather consider the retrieval to be an estimate of a smoothed version of the true state as suggested by Rodgers [2000].

A plot of the error categories is given in Figure 6 and summarized in Table 1. Values in Table 1 are representative of the 24–32 km, the portion of the BrO profiles that are the most robust. The largest error source was the retrieval noise with values near 25%. Other sources were in the 5–10% range. To determine the total error, a distinction between random and systematic sources of error was made. Model errors and model parameter errors are generally considered systematic errors, including the error in BrO cross section. An exception to this is the slit width, or resolution, error. As described in section 3.7, when OSIRIS temperatures depart from nominal values, the slit width, or resolution, is affected. A robust correction is applied such that retrievals, in a monthly mean sense, make use of the slit width appropriate for that monthly mean instrument temperature. Thus, any remaining slit width error, estimated to be 5%, can be considered random. Adding the random errors in quadrature gave a total random error, or precision, of about 27%. The simple sum of the systematic errors, or bias, is 17%. For individual BrO profiles, random and systematic errors may be added in quadrature for a total error of 32%. It is reiterated here that as a result of not being able to quantify the smoothing error, this error budget is appropriate for a smoothed version of the true state. When monthly means of BrO are considered, the random error component becomes

![Figure 6](image-url)
Table 1. Error Budget

<table>
<thead>
<tr>
<th>Error Source</th>
<th>Type</th>
<th>Magnitude (%)</th>
</tr>
</thead>
<tbody>
<tr>
<td>Retrieval noise</td>
<td>R</td>
<td>25</td>
</tr>
<tr>
<td>Forward model parameter</td>
<td></td>
<td></td>
</tr>
<tr>
<td>Resolution</td>
<td>R</td>
<td>10</td>
</tr>
<tr>
<td>All others</td>
<td>S</td>
<td>6</td>
</tr>
<tr>
<td>Forward model</td>
<td>S</td>
<td>3</td>
</tr>
<tr>
<td>BrO cross section</td>
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</tr>
<tr>
<td>Total BrO systematic</td>
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<td>17</td>
</tr>
<tr>
<td>Total BrO random</td>
<td></td>
<td>27</td>
</tr>
<tr>
<td>Inferred precision, Figure 11</td>
<td></td>
<td>~30</td>
</tr>
<tr>
<td>Total BrO Error</td>
<td>S</td>
<td>32</td>
</tr>
<tr>
<td>Photochemistry</td>
<td>S</td>
<td>7</td>
</tr>
<tr>
<td>BrO systematic error</td>
<td>S</td>
<td>24</td>
</tr>
</tbody>
</table>

*Error type is either random (R) or systematic (S). Random errors were added in quadrature. The total error in BrO was obtained by adding the total random and total systematic errors in quadrature.

small and the systematic component becomes the dominant source.

[33] As mentioned in section 3.3, photochemical box model calculations were used to map the OSIRIS BrO to the local time of another instrument to facilitate a comparison or to derive total BrO. In each application this introduces an additional source of systematic error, called photochemical error. Photochemical error includes errors in the box model constraints such as the prescribed ozone, temperature, or families, and rate constants. Sioris et al. [2006] found that the key components in the photochemical error are in the assumed values of air density, NO2, and the rate coefficients of the primary formation and destruction reactions,

\[(R3) \quad \text{BrONO}_2 + h\nu \rightarrow \text{BrO} + \text{NO}_2\]

\[(R4) \quad \text{BrO} + \text{NO}_2 + M \rightarrow \text{BrONO}_2 + M\]

Following Sioris et al. [2006], the photochemical error was assessed by successively varying air density, NO2, and the reaction rate coefficients by their uncertainty, and then recomputing BrO in the box model for a tropical and mid-latitude scenario. The photochemical error is the sum of these, added in quadrature. An error of 7% is representative of altitudes above 20 km for OSIRIS AM BrO and all altitudes for OSIRIS PM BrO. Below 20 km, AM errors were found to be larger, up to 40%, due to uncertainties in the photolysis rate coefficient of reaction (R3). The abundance of BrO is very sensitive to the rate of release from BrONO2 following its buildup throughout the night. Photochemical error is treated as a systematic source.

3.7. Observations During Eclipse Season

[34] Every year between May and July, Odin spends a significant fraction of its orbit in eclipse, with this fraction peaking in June. This results in a gradual cooling of the spacecraft. Sensors in the OSIRIS optics indicate that the temperature drops from a nominal value of 22°C to a minimum of about 10°C. This acts to (1) induce a small wavelength shift, (2) blur the image as manifested by an increase in the slit function width, and (3) introduce a TH offset [McLinden et al., 2007]. This last effect is brought about by a thermal flexing of the spacecraft which alters the alignment of OSIRIS relative to the star tracker, from which OSIRIS obtains its pointing information. Each effect represents a potential source of additional error in the BrO retrievals. The 2001–2008 time series of monthly mean optics temperatures is shown in Figure 7a. While the eclipse period lasts from May–July, there appear to be some lingering effects through to September.

[35] For each zonal mean scan the spectral fitting algorithm determines any relative wavelength shift between the OSIRIS spectra and the cross-section data. During eclipse periods, OSIRIS wavelengths were found to be blue-shifted by about 0.1 nm, corresponding to about one quarter of an OSIRIS pixel. However, the calculated shift was applied to the OSIRIS spectra and thus eliminated the wavelength shift as a potential error source.

[36] Prelaunch testing suggested that a decrease in the OSIRIS optics temperature from 22°C to 10°C leads to a 20% increase in the slit width. As seen in Figure 6, a 5% increase in the slit width leads to an underestimation of BrO in the neighborhood of 10%. In other OSIRIS data products such as NO2 this effect was much smaller, about 3% [Haley et al., 2004], and so was ignored. The reason for the increased sensitivity in BrO is primarily from the interference of O3 in the spectral fitting. To account for this, the cross-section data (and pseudo-absorber basis functions) were generated for a set of slit width scalings: 0.9, 0.95, 0.975, 1.0, 1.025, 1.05, 1.1, 1.15, 1.2, 1.3 and 1.4, and an expression relating optics temperature, \(T_{\text{opt}}\), to relative slit width, \(w\), based on the prelaunch test data was derived,

\[w = 0.8936 + 0.1064 \exp[-0.0819(T_{\text{opt}} - 22)]\]

(1)

where \(T_{\text{opt}}\) is in °C and gives a value for \(w\) of 1 for \(T_{\text{opt}} = 22°C\) and 1.18 for 10°C. On the basis of the optics temperature for a given scan, the cross sections calculated for the nearest slit width scaling were used in the inversion. Values of \(w\) for monthly mean optics temperatures are shown in Figure 7b. The standard (scaling of 1) cross sections were generally used for data from October–April. Data from May–August use reduced resolution cross sections with September somewhere between. Future data versions may employ an alternative approach in which the relative slit width is calculated from measured spectra.

[37] The final eclipse effect, a TH offset of up to 500 m, has been observed in comparisons between OSIRIS ozone and ozonesondes [McLinden et al., 2007]. This leads to a systematic error in the retrieved BrO of 0–6% as seen in Figure 6 with opposite signs at the top and bottom of the profile. Given its relatively small error contribution, no correction was applied for this effect in the current data version.

4. BrO Climatology

[38] OSIRIS BrO monthly means have been calculated using the full 7+ years of data from November 2001 to December 2008. Averages were computed using the same altitude, latitude, and AM/PM binning as the zonal mean L1 spectra. Individual data points with a response of 0.67 or smaller or a relative error of 100% or greater have been excluded. A response threshold of 0.67 was chosen to ensure data points that are heavily influenced by a priori information are excluded [Urban et al., 2005]. Because of
the steep gradients near sunrise and sunset, BrO is particularly susceptible to diurnal effect errors, or the diurnal variation in BrO along the line of sight due to a changing local SZA [McLinden et al., 2006]. To minimize the impact of this error source, only SZAs of 87° or smaller were included in the monthly means. Finally, only mean values made up of at least 20 individual data points were considered. After this filtering, monthly mean profiles typically extend from 16 to 34 km in the tropics and 18–36 km in the extratropics.

Figure 8 shows the resulting monthly mean AM (descending node) BrO (BrO-AM) maps. There is a minimum of 3 pptv at the tropical tropopause and an increase with altitude and poleward to a maximum of about 16 pptv. Some seasonal features can be understood by examining OSIRIS monthly mean NO$_2$ number density [Brohede et al., 2007b], shown in Figure 9 (also calculated using SZAs of 87° or less).

[40] Because of the coupling of the nitrogen and bromine families via reactions (R4) and (R3) there is expected to be a relationship between BrO and NO$_2$. Months such as March and October show a local maximum in BrO around 24 km in the north polar region that can be attributed to the minimum in NO$_2$ due to its conversion to NO$_3$ reservoir species, thereby making it unavailable to form BrONO$_2$. (It is expected that a similar feature would be present through the NH winter.) Likewise, in the NH summer, the BrO contours are much wider indicating reduced values. Again this is consistent with the NO$_2$ which shows a maximum in this region. An overall seasonal cycle is also evident with a maximum in the winter pole and a minimum in the summer pole. This is related to the seasonal cycle in NO$_2$ and is consistent with the findings of Pundt et al. [2002].

[41] A BrO-NO$_2$ correlation plot is shown in Figure 10 at a latitude of 65°N where there is a large seasonal cycle and 8 months of coverage. Correlation coefficients in the lower stratosphere (16–26 km) were typically −0.9 where the coupling is expected to be tight, but above this values were small and slightly positive. At 34 km, for example, BrO is completely unaffected by the NO$_2$ seasonality. The reduced sensitivity to NO$_2$ above 30 km is a result of the rapid decrease in the production rate of BrONO$_2$, reaction (R4), from the declining densities of NO$_2$ and air.

[42] The monthly BrO AM standard deviations, expressed as a fraction of the mean, are shown in Figure 11. Values between 0.2–0.35 dominate, with occasional pockets up to 0.5. The larger standard deviation at the South pole from September to November below 20 km is likely due to differing amounts of denitrification from year to year. The other pockets of higher standard deviation tend to coincide with smaller absolute amounts of BrO. In general, sources of
Figure 8. OSIRIS monthly mean descending node (AM) BrO mixing ratio (in pptv), averaged over November 2001 to December 2008.

Figure 9. OSIRIS monthly mean descending node (AM) NO$_2$ number density, averaged over November 2001 to December 2008.
variability include natural day-to-day, interannual variability, a drifting of the Odin orbit in LST, and noise in the retrieval. The contribution from Odin’s drift in LST can be assessed by examining the monthly standard deviations of the a priori profiles. Removing this from the retrieved BrO standard deviations leads to a 3–5% reduction. Considering all this, the precision (random error) of a typical daily profile can be considered to be ~30%. This can be contrasted with the theoretical precision estimate of 27% from Table 1.

Since Odin possesses a sunrise-sunset orbit, OSIRIS has the advantage of measuring at two local times (~0630 and 1830), although because of the precession of Odin’s orbit toward later local times, the AM half of the orbit is favored with little or no PM coverage in the tropics from about 2004 onward. Coverage is further limited by the SZA limit of 87°. As a check of the internal consistency of the OSIRIS BrO product, AM-PM monthly mean differences were calculated and compared to monthly mean a priori differences (not shown). There is a consistent month-to-month picture, with generally larger mixing ratios in the AM for northern extratropics but more of a mixture in the southern extratropics. Since SZAs are smaller in the northern hemisphere summer than the southern hemisphere summer, this hemispheric difference is not necessarily unexpected.

5. Comparisons With Other Instruments
5.1. Diurnal Mapping

Because of its diurnal nature, comparisons between instruments measuring BrO at different local times must be made with caution. As was done for the validation of OSIRIS NO2 [Brohede et al., 2007a] as well as ACE-FTS NO and NO2 [Kerzenmacher et al., 2008], in this work OSIRIS BrO was mapped from its local time to that of the other instruments.
This was done using box model calculations using the following expression:

$$n_{OS}(LST_i) = n_{OS}(LST_{OS}) \left( \frac{n(LST_i)}{n(LST_{OS})} \right)_{model}$$

(2)

where \( n \) is the BrO number density and the subscripts “OS”, “\( i \)”, and “model” represent OSIRIS, other instrument, and box model, respectively. For single-profile comparisons the box model is constrained with ECMWF temperature and simultaneously measured OSIRIS ozone [Haley and Brohede, 2007], averaged over the same scans that were used in deriving the zonal L1 data, as well as monthly mean Odin NO\(_2\) [Brohede et al., 2008]. When monthly means were compared, the box model was initialized with monthly mean ECMWF temperature, OSIRIS ozone, and Odin NO\(_2\). This assumes the partitioning of Br\(_y\) among its family members is largely independent of the absolute amount of Br\(_y\) specified.

5.2. Single Profile Comparisons

The zonal mean nature of OSIRIS BrO means it is difficult to derive quantitative information when comparing to single BrO profiles. Nonetheless, it is still worthwhile given the limited sources of correlative BrO profiles. OSIRIS BrO was compared with profiles measured from the SAOZ–BrO balloon-borne instrument [Pundt et al., 2002], a UV spectrometer designed specifically to measure BrO via solar occultation during the ascent of the balloon. Comparisons were only considered if OSIRIS measurements exist for the same day, within 5° latitude of a SAOZ profile, and if the individual scans contributing to the zonal L1 spectra bracketed the SAOZ longitude. Six OSIRIS–SAOZ profiles matched these criteria, and they are shown in Figure 12. The common altitude range is 16–30 km. These SAOZ profiles were recorded in the late afternoon between SZAs of 83–87°. Both the original OSIRIS profile and the OSIRIS profile mapped to the SZA of SAOZ are shown. Overall there is a general consistency between the two profiles, with some larger differences where the error bars do not overlap, particular for the Niamey comparison. However, it is difficult to differentiate between actual inconsistencies and those differences caused by the instruments viewing differing air masses. All mapped OSIRIS profiles agreed better with SAOZ than did the originals.

5.3. Monthly Mean Profile Comparisons

Comparisons of monthly mean OSIRIS profiles were made with those from SCIAMACHY, a satellite
instrument on-board ENVISAT [Bovensmann et al., 1999]. SCIAMACHY, like OSIRIS, measures BrO profiles using limb-scattered sunlight [Sinnhuber et al., 2005]. The relative difference between OSIRIS BrO-AM and SCIAMACHY monthly means, with OSIRIS interpolated to the SCIAMACHY grid and mapped to the LST of ENVISAT (~1000), is shown in Figure 13 over their common retrieval range, 16–32 km. Latitudes poleward of 70° were excluded as the SCIAMACHY BrO in these regions may be subject to large diurnal effect errors. The relative difference was calculated using (OSIRIS - SCIAMACHY)/SCIAMACHY. All months tended to display a similar pattern with OSIRIS larger in the topics above about 26 km. Poleward and below this SCIAMACHY BrO tends to be larger. Contrasting the differences in the NH between June and October, it is clear that OSIRIS possesses a stronger seasonal cycle than SCIAMACHY. This could be attributed to an incomplete removal of the eclipse effect, but if that were the case it should also be visible at tropical latitudes, which is not. In the lowest OSIRIS retrieval altitudes, where there may be some edge effects. The overall level of agreement between the two monthly mean data sets is in the 30–40% range. This can be compared with the estimate of OSIRIS systematic errors, about 25% including a photochemical contribution.

A comparison with profiles derived from ground-based zenith-sky UV-visible spectrometers at Harestua, Norway (60°N), and Observatoire de Haute-Provence (OHP), France (44°N), has also been conducted. Details of the retrieval algorithm are available elsewhere [Hendrick et al., 2008], but in short, it uses a combination of spectral fitting to obtain SCDs and optimal estimation to invert the SCDs to number density profiles in conjunction with a radiative transfer model and a photochemical box model. The resolution of the profiles is rather coarse at ~10 km with roughly two independent pieces of information in the stratosphere [Hendrick et al., 2007]. The Harestua data set used in this study spans 2002–2006, very close to the OSIRIS coverage. OHP is shorter, 2005–2006. Monthly mean Harestua total VCDs were found to be in good agreement with SCIAMACHY and Global Ozone Monitoring Experiment (GOME) VCDs. The comparison of OSIRIS monthly means (average of 55° and 65°N) with the Harestua monthly means is shown in Figure 14 for the 8 months of OSIRIS coverage at this latitude. OSIRIS profiles have been mapped to a SZA of 80° and then convolved with the Hendrick et al. [2007] averaging kernels. The comparison was generally consistent with that of SCIAMACHY in that in the summer months, OSIRIS was consistently smaller, by up to 25%, but for March and October, there is excellent agreement. The comparison with OHP was very similar (not shown).

6. Implication for Total Br$_y$

A number of recent studies have attempted to quantify the total amount of inorganic bromine in the stratosphere, and hence the contribution from VSLS. These include ground-based UV-vis instruments [Hendrick et al., 2007], the balloon-borne DOAS/LPMA [Dorf et al., 2006, 2008], MLS/Aura [Kovalenko et al., 2007], SCIAMACHY from two different research groups [Sinnhuber et al., 2005; Sioris et al., 2006], and a balloon-borne whole-air sampler [Laube et al., 2008]. The Laube et al. [2008] measurements...
differ from the others in that bromocarbons are measured in situ throughout the upper troposphere and stratosphere. These estimates vary from 17.5 to 25 pptv and are based on measurements made after 2000, and hence probe air entering the stratosphere after ~1995. Between 1995 and 2005 the sum of CH$_3$Br and the halons delivered into the stratosphere has been roughly constant at 15.5–17 pptv due to an off-setting decline in CH$_3$Br and increase in halons, with a peak in 1998 [WMO, 2007]. As such, to a first approximation, it is fair to compare these results with OSIRIS without regard to year of measurement or age of air. If the contribution of CH$_3$Br and halons to total Br$_y$ is taken at 16 pptv over this period, then this would imply a VSLS source of 1.5 to 9 pptv. This can be compared with the Laube et al. [2008] measurements of organic bromocarbons in the tropical tropopause layer which suggested a contribution from VSLS of 1.3 pptv. These varied results highlight the uncertainty in the VSLS contribution to Br$_y$.

[49] In this section, two methods were used to obtain an estimate of total stratospheric Br$_y$. The first involves comparing OSIRIS partial VCDs to those predicted by the photochemical box model initialized with different amounts of VSLS. A similar approach was used by Salawitch et al. [2005] and Sioris et al. [2006]. The box model was run for the 15th of each month and constrained with monthly mean ECMWF temperature, OSIRIS ozone, SMR N$_2$O, and Odin NO$_y$. VCDs were calculated but restricted to the common OSIRIS retrieval range over all latitudes: 20–34 km. This also ensured that only the highest retrieval responses were used. All scenarios used the Wamsley “organic” Br$_y$ expression, updated to 2005 (the midpoint of the OSIRIS data set), but with a VSLS contribution, constant with altitude, such that the total Br$_y$ is varied from 19 to 25 pptv in 1 pptv increments. Adding a constant amount of Br$_y$ at all altitudes assumes the source is short-lived and releases its bromine (or, converted to inorganic form) in the lowermost stratosphere. It also assumes that there is no trend in the source since it is constant with age of air.

[50] Results for the month of September are shown in Figure 15a. Calculating the minimum RMS differences over latitude for the different scenarios suggested the best agreement was for total Br$_y$ of 21–22 pptv. This value was further refined by fitting a quadratic to the three smallest RMS values and finding the Br$_y$ amount that minimized the RMS quadratic. For September this gave a values of 21.1 pptv. Analogous values for all months are shown in

![Figure 14.](image-url)
Figure 15. (a) Comparison of monthly mean OSIRIS BrO (AM) partial vertical column densities (VCDs) (19–35 km) to those from a photochemical box model assuming varying levels of Br\textsubscript{y} contribution from very short lived substances (VSLS) for September. Solid lines show model calculated partial VCDs for a total Br\textsubscript{y} of 19–25 pptv (19 and 25 pptv values are indicated with blue lines; other colors represent intermediate values in 1 pptv increments), and squares indicate OSIRIS values. Considering all latitudes, a total Br\textsubscript{y} of 21.1 pptv best agrees with the OSIRIS partial VCDs. (b) Inferred total Br\textsubscript{y} for each month. The horizontal line indicates the mean over all months, (21.2 ± 0.5) pptv.

The inferred OSIRIS Br\textsubscript{y} from the correlation plot, representative of 2005, is 21.0 pptv, chosen as it is the maximum value (outside of the data point corresponding to 36 km, or low N\textsubscript{2}O, which may be contaminated by a

\[ n_{OSIRIS} = n_{OS} / \frac{n_{BrO}}{n_{model}} \]  

analogous to the method employed by Kovalenko et al. [2007] and numerous other studies. The Br\textsubscript{O}/Br\textsubscript{y} ratio was calculated using the photochemical model set to the midpoint of each month and constrained with monthly mean OSIRIS \textrm{O}_{3}, ECMWF temperature, SMR N\textsubscript{2}O, and Odin NO\textsubscript{y}. Once an initial estimate of the Br\textsubscript{y} was obtained, it was then used in the box model to recalculate Br\textsubscript{O}/Br\textsubscript{y} in case the partitioning was sensitive to the assumed Br\textsubscript{y} profile. Monthly Br\textsubscript{y} calculated in this fashion from the Br\textsubscript{O}-AM data are shown in Figure 16. The morphology is similar to Br\textsubscript{O} but with reduced seasonality. Some seasonality is expected given the Brewer-Dobson circulation pattern. There is a plateau of 18–22 pptv above 28 km or so, with some months displaying the maximum in the tropics, others in the extratropics. The extratropics also show a small decline above the peak which may be an age-of-air effect. The uppermost retrieval level often shows a 1–2 pptv jump, similar to that seen in Br\textsubscript{O}, which appears to be an artifact of the retrieval related to the amount of Br\textsubscript{O} assumed above the top of the retrieval range. In some instances, too little Br\textsubscript{O} may be assumed above the 36 km, and thus the Br\textsubscript{O} in the uppermost retrieval level is forced to compensate.

[52] OSIRIS monthly Br\textsubscript{y}, from Figure 16, was examined in tracer space in Figure 17. This shows the correlation between monthly mean SMR N\textsubscript{2}O and OSIRIS Br\textsubscript{y}, considering latitudes between 40°S to 40°N. Tracer space offers the advantage of reducing differences due to age of air. Figure 17a shows the individual OSIRIS monthly Br\textsubscript{y} data points for both AM and PM as well as their values averaged over 20 ppbv-wide N\textsubscript{2}O bins. The values at the smallest N\textsubscript{2}O occur at altitudes of 36 km, and these may be subject to some top of retrieval range effects. Overall the AM and PM are generally similar, and the differences that exist are likely related to the difference in latitudinal sampling with most of the PM coverage originating from the 25° and 35° latitude bins. Figure 17b shows the bin-averaged N\textsubscript{2}O-Br\textsubscript{y} correlation along with the Wamsley et al. [1998] “organic” relationship, which considers only CH\textsubscript{3}Br and halons, adjusted to 2005 as follows. In the tropics the mean age of air over the OSIRIS retrieval range is ∼3 years. Hence the Br\textsubscript{y} maximum in the Wamsley et al. [1998] relationship is scaled to 16 ppbv; a value obtained from WMO [2007] for 2002, the year of stratospheric entry. In Figure 17b the mean Br\textsubscript{y} over all AM and PM monthly values are used. Given the much larger number of AM measurements in this latitude range, about 3 times the PM, the overall mean is heavily weighted toward the AM mean from Figure 17a. Correlations from MLS and a DOAS balloon flight are also shown, taken from Kovalenko et al. [2007], as are total Br\textsubscript{y} estimates from other sources and shown simply as open triangles. The OSIRIS correlation is very similar to the MLS and DOAS curves. By taking the difference between the OSIRIS and Wamsley curves, an estimate of the Br\textsubscript{y} originating from VSLS can be made. This difference was 5 pptv on the low-N\textsubscript{2}O end and 7 pptv on the high end. The simple mean of these values is 5.8 pptv.

[53] The second method of estimating total Br\textsubscript{y} was to derive it directly from Br\textsubscript{O} using photochemical modeling. In this work, Br\textsubscript{y} was derived from the monthly mean Br\textsubscript{O} from section 4 using the expression,
retrieval artifact). This is virtually identical to the 21.2 pptv obtained from the column method. The slow decline with altitude just above this (decreasing N\textsubscript{2}O) appears consistent with an age of air effect. A measure of the uncertainty can be obtained by using the systematic error in Br\textsubscript{y} from Table 1 of 24\%, giving an uncertainty of 5.0 pptv. The relatively constant difference between the OSIRIS and Wamsley curves in Figure 17b down to near-tropospheric values of N\textsubscript{2}O suggests that 5–7 pptv is either released in the lowermost stratosphere and/or is transported across the tropopause already in inorganic form. The lowest tropical altitude contributing to these results is 18 km (17–19 km layer) which corresponds to an age of air of up to 4–5 months. This places an upper limit of ∼4 months on the lifetime of the VSLS candidate species. This would not seem to contradict any current views as, of the VSLS proposed by \textit{WMO} [2007], only CH\textsubscript{3}BrCl exceeds this with a lifetime of 5 months with an estimated mixing ratio of only 0.3 pptv in the tropical tropopause layer [\textit{WMO}, 2007]. The single largest contributor is thought to be CH\textsubscript{2}Br\textsubscript{2} which has a lifetime of ∼4 months and contributes ∼1.8 pptv to inorganic bromine (0.9 pptv mixing ratio) [\textit{WMO}, 2007].

7. Summary and Conclusions

[54] A 7+ year (2001–2008) data set of stratospheric BrO profiles measured by the Optical Spectrograph and InfraRed Imager System (OSIRIS) satellite instrument has been presented. Zonal mean radiance spectra were computed for each day and inverted to yield effective daily zonal mean BrO profiles from 16–36 km. Single profile precision was found to be about 30% with an effective resolution of 3–5 km, respectively, throughout much of the retrieval range. The systematic error is estimated to be about 17%. It was necessary to implement a temperature-dependent slit function width to allow for a blurring of the image due to a cooling during eclipse season. Comparisons between individual profiles and monthly means are found to agree to typically 30% with other observations of BrO from ground-based, balloon, and satellite instruments. However, it is noted that OSIRIS displayed a larger seasonal cycle than some of the correlative data.

[55] A BrO climatology was presented, and its abundance, morphology, and correlation with NO\textsubscript{2} were determined to be consistent with the current understanding of bromine chemistry. Monthly mean BrO, in concert with photochemical modeling, were used to derive monthly Br\textsubscript{y} maps. Two methods of calculating total stratospheric Br\textsubscript{y}, one comparing partial columns of BrO to model simulations, the other involving correlations between Br\textsubscript{y} and N\textsubscript{2}O, presented a consistent picture and suggest (21.0 ± 5.0) pptv with a VSLS contribution of (5.0 ± 5.0) pptv. The Br\textsubscript{y}-N\textsubscript{2}O correlation also placed an upper limit of the lifetime of VSLS species of about 4 months. OSIRIS v4.0 BrO data are available for download in day-based HDFEOS-5 format from http://osirus.usask.ca.
Figure 17. (a) Tracer correlation plot of OSIRIS Br\textsubscript{y} and Odin/Submillimetre and Millimetre Radiometer N\textsubscript{2}O monthly mean from 40°S to 40°N shown separately for AM and PM measurements. Each symbol represents an individual altitude/latitude. Lines represent mean Br\textsubscript{y} averaged over 20 ppbv wide bins in N\textsubscript{2}O and their 1\textsigma standard deviation (shown dashed for PM values). (b) Tracer correlation of Br\textsubscript{y} and N\textsubscript{2}O for different data sources. The open triangles represent inferred total Br\textsubscript{y} from other studies (without accompanying N\textsubscript{2}O). MLS, Microwave Limb Sounder.
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