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Evidence of gravity waves into the atmosphere during the March 2006 total solar eclipse

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Abstract. This study aims at providing experimental evidence, to support the hypothesis according to which the movement of the moon's shadow sweeping the ozone layer at supersonic speed, during a solar eclipse, creates gravity waves in the atmosphere. An experiment was conducted to study eclipse induced thermal fluctuations in the ozone layer (via measurements of total ozone column, ozone photolysis rates and UV irradiance), the ionosphere (Ionosonde Total Electron Content – ITEC, peak electron density height – hmF2), and the troposphere (temperature, relative humidity), before, during and after the total solar eclipse of 29 March 2006. We found the existence of eclipse induced dominant oscillations in the parameters related to the ozone layer and the ionosphere, with periods ranging between 30–40 min. Cross-spectrum analyses resulted to statistically significant square coherences between the observed oscillations, strengthening thermal stratospheric ozone forcing as the main mechanism for GWs. Additional support for a source below the ionosphere was provided by the amplitude of the oscillations in the ionospheric electron density, which increased upwards from 160 to 220 km height. Even though similar oscillations were shown in surface temperature and relative humidity data, no clear evidence for tropospheric influence could be derived from this study, due to the modest amplitude of these waves and the manifold rationale inside the boundary layer.

1 Introduction

Gravity waves (GWs) constitute an essential motion component of the atmospheric circulation due to their important contribution in the momentum and energy budget and in the wind systems of the atmosphere (Fritts and Alexander, 2003). Their role in weather, climate and atmospheric chemistry is very important. In particular, they can transport energy and momentum between different atmospheric regions and disturb the balanced state or initiate and modulate convection and subsequent hydrological processes (e.g. Mapes, 1993). The initiated by GWs convection, can result to formation of clouds (e.g. orographic cirrus, polar stratospheric and mesospheric; Dörnbrack et al., 2002), modify chemistry, and trigger numerous processes (e.g. Voigt et al., 2000). During their dissipation phase, they contribute to the vertical transport and mixing of chemical species and influence momentum and energy in the upper troposphere/lower stratosphere and in the mesosphere and lower thermosphere (e.g. Hays et al., 2003).

A varying number of sources are believed to generate GWs at lower levels of the atmosphere including topography, convective and frontal activity, wind shear and geostrophic adjustment, while at greater heights their sources include non-linear wave-wave interactions, auroral currents, ion drag and Joule heating but also the differential heating of the atmosphere at dawn and dusk terminator and during solar eclipses (Fritts and Luo, 1993 and references therein).

Chimonas and Hines (1970) were the first to suggest that during a solar eclipse the disturbance of the heat balance along the supersonic travel of the trajectory of the moon's shadow could generate GWs. The source of these waves was assumed to be either at higher altitudes e.g. at around 90 km

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where molecular oxygen heating begins (Chimonas, 1970), or at lower altitude e.g. water vapor IR absorption (Lamb waves) and the ground cooling (Chimonas and Hines, 1971; Chimonas, 1973). Moreover, in both Chimonas (1970) and its subsequent follow on study of Firtts and Luo (1993), thermal cooling of the stratospheric ozone layer as the forcing function of GWs during a solar eclipse was considered. Eckermann et al. (2007), with their high-altitude global numerical weather prediction model, recently simulated an induced radiative cooling rate in the stratosphere, 2–3 times larger than assumed in these earlier stratospheric bow wave models. In all cases, the source ambiguity makes subsequent experimental studies seeking evidence of eclipse-generated gravity waves, particularly complicated.

There have been a number of attempts to detect ground level atmospheric pressure waves resulting from solar eclipses. The periods of such waves range from the order of 1 min to 1 h or so, the decay times range from half the wave period upwards, and the ground level amplitude is unlikely to exceed 0.1 hPa, and is typically far less (Jones, 1999). In only few cases does the balance probability lie in favour of a detection of such waves (Jones, 1999; references therein). Anderson et al. (1972) reported on surface pressure fluctuations in the range 15–150 min and Seykora et al. (1985) found a surface pressure response with a period of about 4 h and a speed of 320 m s^{-1} . The amplitudes of these waves were larger than those predicted by Chimonas (1970), but agree well with the 0.1–0.5 hPa range modelled recently by Eckermann et al. (2007).

A limited number of attempts to detect GW signals on total ozone have been made. Mims and Mims (1993) identified a sequence of 4–5 nearly uniformly spaced fluctuations, however with small periods (4.5–7.2 min), on total ozone series, using a TOPS (Total Ozone Portable Spectrophotometer), but the linkage with eclipse induced GWs appears weak. Zerefos et al. (2000), deploying power spectral analysis on erythral irradiance, revealed a significant oscillation in the ozone layer with a period of about 20 min.

In the ionosphere, waves with a period of 10–40 min have been measured at more than 500 km from the zone of totality (Singh et al., 1989), traveling at subsonic (Davis and da Rosa, 1970) or supersonic velocities (Hanuise et al., 1982). Lastly, a source location has been identified in the thermosphere at 170 km altitude due to reduced heating by absorption of extreme ultraviolet solar radiation, based on the analysis of ionosonde measurements (Liu et al., 1998; Altadill et al., 2001; Sauli et al., 2006), while larger periods of about 1 h have also been reported (Altadill et al., 2001). Prior to those observations, the idea of an in-situ thermospheric wave source was highlighted by various modeling studies (Ridley et al., 1984; Roble et al., 1986; Müller-Wodarg et al., 1998).

Despite decades of research, observational evidence for a characteristic bow-wave response of the atmosphere to eclipse passages remains equivocal (Eckermann et al., 2007). In this work, an attempt to shed light on the generation of

GWs during solar eclipses is made, with measurements at the three critical layers in the atmosphere namely the troposphere, the stratosphere and the ionosphere. The main goal is to provide experimental evidence supporting the initial hypothesis that the cooling of the ozone layer in the stratosphere by the moon shadow travelling at supersonic speed during the eclipse, constitutes a source of gravity waves propagating both upwards and downwards.

2 Instrumentation

Ultraviolet direct and global solar spectral measurements were performed at Kastelorizo (36°09' N, 29°35' E), Thessaloniki (40°38' N, 22°57' E) and Athens (38°03' N, 23°52' E) using Brewer spectroradiometers (MKIII, MKII and MKIV, respectively). A map with the location of the above stations with regard to eclipse path and circumstances can be found in Gerasopoulos et al. (2007)¹. The characteristics of the instrument and details about their calibration procedures are described in detail by Bais et al. (1996). The three instruments followed the same measuring schedule on the eclipse day (29 March) and on the previous day. For this study the operating software of the instrument was modified to allow the alternating measurement (every 30 s) of global and direct spectral irradiance at 6 wavelengths between 302 and 320 nm that are used for measuring total column ozone and columnar SO₂. These measurements were performed from 07:30 UTC until 14:30 UTC. More details are given by Blumthaler et al. (2006) and Kazadzis et al. (2007).

Ozone photolysis frequencies (JO^{1D}) have been also calculated from the Brewer spectroradiometer global irradiance measurements following the methodology described in Kazadzis et al. (2004). JO^{1D} measurements have been also conducted at Finokalia, with a filter radiometer – Meteorologie Consult, Germany (Gerasopoulos et al., 2006). We have also used UV measurements from the Greek UV monitoring network (<http://www.uvnet.gr>), equipped with NILU-UV multi-channel radiometers, providing UV irradiance at five wavelength bands centered at 305, 312, 320, 340 and 380 nm (Kazantzidis et al., 2007). Meteorological measurements at a number of sites including Kastelorizo and Finokalia have been also conducted during the eclipse, as described by Founda et al. (2007).

Ionospheric observations from the National Observatory of Athens Digisonde (<http://www.iono.noa.gr>) were used for the investigation of the ionospheric response to the solar eclipse of 29 March 2006 over Athens (38°00' N, 23°30' E). In particular, calculations of electron density profiles up to 1000 km, obtained under a vertical incidence ionospheric sounding campaign, were used for the derivation of the peak electron density height, hmF2, and the estimation of the

¹Gerasopoulos, E., Zerefos, C. S., Tsagouri, I., et al.: The Total Solar Eclipse of March 2006: Overview, *Atmos. Chem. Phys. Discuss.*, in preparation, 2007.

Ionosonde Total Electron Density, ITEC., (Gerasopoulos et al., 2007¹). In addition, the electron densities at fixed ionospheric altitudes zones were calculated, enabling the study of electron density variations as a function of time and altitude (Altadill et al., 2001).

3 Analysis and results

3.1 Methodology

Observations from various atmospheric layers were analyzed deploying Spectral Fourier Analysis, in an effort to investigate the possible detection of GWs propagating in different atmospheric heights during the 29 March 2006 Total Solar Eclipse. In this study we have used the following parameters: i. total ozone over Kastelorizo, Athens and Thessaloniki, mainly as an index of the disturbance in the stratospheric ozone layer, additionally supported by ozone photolysis rates JO^1D and UV (305 nm) irradiance measurements, ii. the Ionosonde Total Electron Content (ITEC) as an indicator for the total ionospheric ionization disturbances, and the peak density height of the F2 layer (hmF2) as a first indicator for the propagation of wave-like motions in the ionosphere, iii. ground measurements of temperature and relative humidity at Kastelorizo and Finokalia.

As input data for the spectral analysis, we have used the residual time series of the above parameters. Residuals have been calculated by applying polynomial fittings to the initial time series, to remove the combined effect of the eclipse and the diurnal variability of each parameter. We have chosen the maximum period during the eclipse that almost cloud free conditions were encountered at all sites, so that the same period would be used for all parameters (10:00–11:52 UTC). Short-time contamination by light and temporary cloudiness or a very few cases of aircraft contrails (evidenced mainly on actinometric data) at each site, has been also removed by substituting the respective residual values with zeroes. For all parameters except for the ionospheric, 1 min time resolution data were available. Ionospheric parameters were available with a 4 min time resolution.

Because of the short duration of the investigated period, we have used the zero-padding method which consists in adding zeros at the end of the time series, to improve frequency resolution of the resulting power spectra, thus enhancing our ability to estimate the signal's period more accurately and discriminate close periodicities.

3.2 Investigation of observed oscillations in the ozone layer

For the period between 10:00 and 13:00 UTC, the total ozone column at Kastelorizo was changing on the day before the eclipse from about 335 DU down to about 325 DU, and on the day of the eclipse it was increasing from about 290 DU to 305 DU (Kazadzis et al., 2007). An increase of total ozone

by the end of the eclipse in the range 5–24 DU is also reconfirmed from the NILU-UV measurements in a number of stations (Kazantzidis et al., 2007). During the course of the eclipse, a gradual decrease in total ozone at Kastelorizo, followed by an almost symmetric increase after the totality is observed (Fig. 1a, upper graph). This effect has been reported earlier by Zerefos et al. (2000), and it was attributed partly to the limb darkening effect and partly to the increasing influence of direct irradiance by the diffuse radiance entering the field of view of the total ozone instruments. Blumthaler et al. (2006), confirmed this finding by calculating a correction for total ozone due to limb darkening. This correction was found to be very small (less than 0.01%). Therefore the most possible reason for the reduced total ozone values during the eclipse is the contamination of direct irradiance measurements by the diffuse radiation (Kazadzis et al. 2007).

Total ozone residuals at Kastelorizo are shown in Fig. 1a (lower graph). A Savitzky-Golay smoothing (2nd order moving polynomial, 10 points) is also used for the better visualization of the lower frequency fluctuations. The two vertical lines in Fig. 1a correspond to sun coverage by the moon of more than 70%, corresponding roughly to a reduction in direct irradiance measurements (under normal conditions) at air mass factors of more than 3, which are usually discarded in the standard Brewer total ozone measurements (Kazadzis et al., 2007). This effect results in a smooth continuous reduction of the ozone values, which is removed by the applied polynomial fit, thus the eclipse induced oscillations in the total ozone data are added on top of this smooth reduction of total ozone. The peak-to-peak amplitude of the residuals is about 2–3.5% of the total ozone averaged over the same period, and three main oscillations were observed starting 30 min after the first contact. Cloud development a little before last contact prevented the further capture of the evolution of the waves.

Since these oscillations are observed mainly during the measurement interval influenced by the diffuse radiation, we have used two additional parameters to express total ozone variability and confirm the existence of the oscillations, the photolysis rate of the reaction of O_3 to $O(^1D)$, JO^1D and UV irradiance at 305 nm. JO^1D is mainly sensitive to the total column ozone, but also to the total aerosol optical depth, aerosol growth with relative humidity and by clouds depending on the optical depth of the cloud (Ruggaber et al., 1994). Moreover, UV irradiance at 305 nm exhibits maximum ozone absorption compared to higher wavelengths provided by the NILU-UV radiometers (Kazantzidis et al., 2007). It should be noted that the effect of diffuse irradiance that was regarded as the reason for the biased total ozone measurements during the eclipse, has no influence on the Brewer irradiance measurements, and consequently on the derived JO^1D data.

JO^1D measurements at Kastelorizo and Finokalia are shown in Figs. 1b and c (upper graphs) with the respective polynomial fittings. JO^1D residuals are presented in the

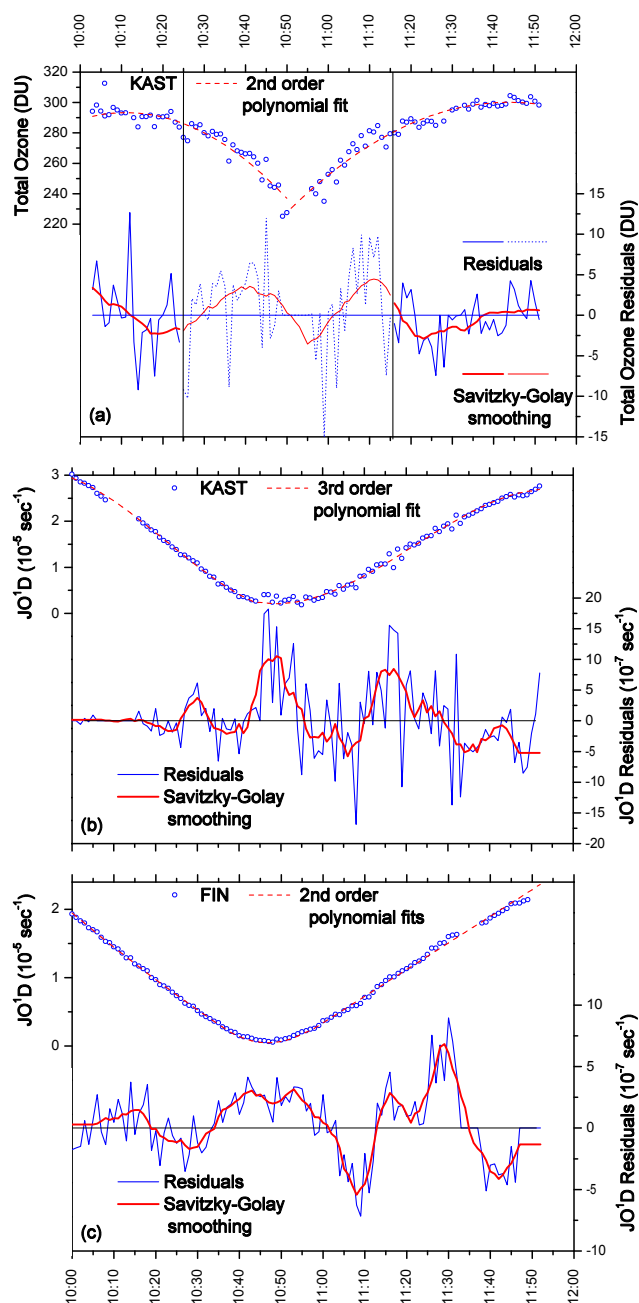


Fig. 1. (a) Upper: total ozone over Kastelorizo during the eclipse and second order polynomials fitted on the data, Lower: total ozone residuals and a Savitzky-Golay (2nd order moving polynomial, 10 points) smoothing line. The two vertical lines correspond to sun coverage by the moon of more than 70% and residuals and smoothing line during this period are presented by dot and thin (red, continuous) lines, respectively, (b) Upper: JO^1D at Kastelorizo (KAST) during the eclipse and a third order polynomial fitted on the data, Lower: JO^1D residuals and a Savitzky-Golay smoothing line, and (c) Upper: JO^1D at Finokalia (FIN) during the eclipse and second order polynomials fitted on the data, Lower: JO^1D residuals and a Savitzky-Golay smoothing line. Time in UTC.

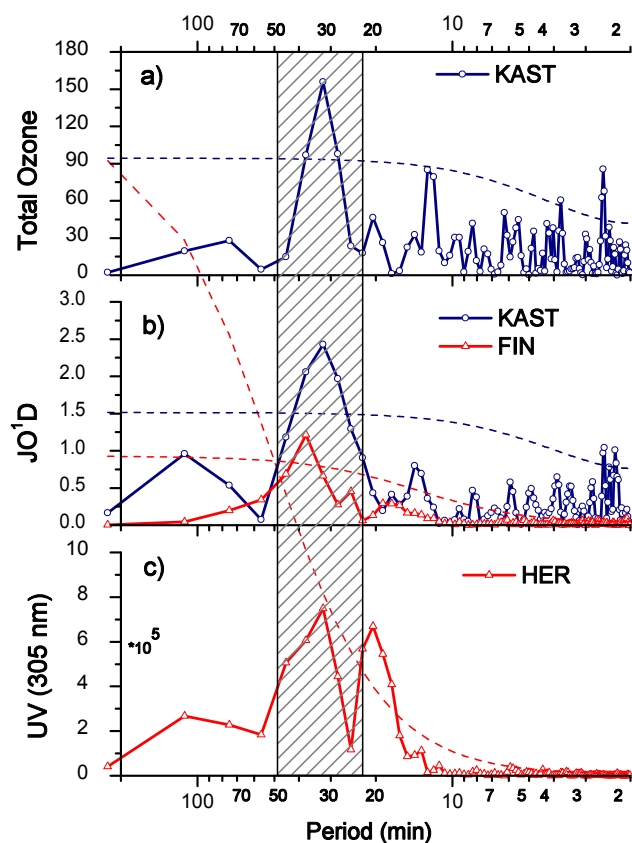


Fig. 2. Power Spectrum Analysis applied on the residuals of various atmospheric parameters during the eclipse (10:00–11:52 UTC, 29 March 2006); Y-axes correspond to the spectral estimates for the following parameters: (a) total ozone at Kastelorizo (KAST), (b) JO^1D at Kastelorizo (blue line + circles) and Finokalia (FIN, red line + triangles), (c) UV (302 nm) at Heraklion (HER, red line + triangles), Lefkosia (LEF, blue line + circles) and Ioannina (IOA, green line + triangles reversed). Similarly colored dash lines represent the 95% confidence limits of each spectrum. Axes in (c) are divided by 10^5 .

lower graph of Figs. 1b and c, and the smoothing line reveals oscillations with similar characteristics. Indeed, the main oscillations found in total ozone at Kastelorizo are reproduced with a certain lag in JO^1D at the same station, but also in distance from totality, at Finokalia. The average peak-to-peak amplitude of the residuals is about 8% and 10% of the JO^1D over the same period, at Kastelorizo and Finokalia, respectively, however corresponding to almost half amplitudes for the more distant from totality station (Finokalia), given the different levels of JO^1D at the two sites.

The oscillations are further investigated with Spectral Fourier Analysis (Fig. 2). The power spectrum of total ozone over Kastelorizo reveals a significant oscillation (99% confidence level, not shown) with a period in the range 28–38 min (Fig. 2a). A secondary oscillation is found at periods 12–13 min which approaches the 95% confidence level.

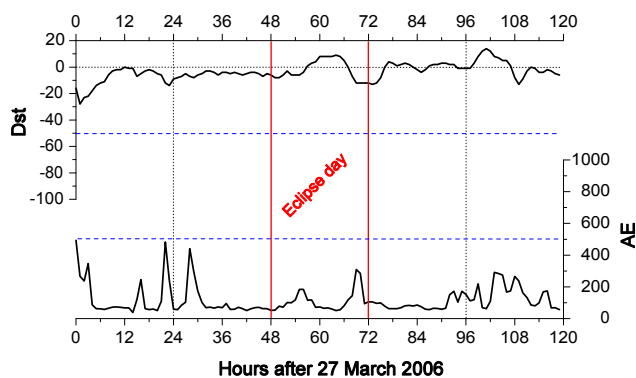


Fig. 3. Hourly records of D_{st} index (<http://swdcwww.kugi.kyoto-u.ac.jp/dstdir/index.html>), as indicator of the geomagnetic activity level (top panel) and of AE index (<http://swdcwww.kugi.kyoto-u.ac.jp/aedir/index.html>) indicating the level of the auroral electrojets intensity (bottom panel) for the time interval 27–31 March 2006. The blue dash lines represent indicative thresholds for the identification of considerable geomagnetic/magnetospheric disturbances.

The same process was applied on the total ozone time series at Thessaloniki (not shown). A significant oscillation (95% confidence level) with a period of 28–32 min was found, denoting the spatial extent of the ozone layer perturbation. The amplitude of the residuals was found to be in the order of 3–4% of the total ozone, averaged over the same period. The influence of clouds was much more extended on total ozone over Athens, thus hindering the attribution of fluctuations to GWs. The discussed range of periodicities could not be identified on the previous and next day spectra of the total ozone series over all sites (for the same time interval; not shown).

The power spectrum of JO^1D at Kastelorizo reveals the same features with total ozone at the same station, allowing its use as a proxy for total ozone variability. The same spectrum is also revealed when 4-min averages are first extracted (not shown). The latter ensures that no bias is imposed on the cross-spectrum between the 4 min averages of total ozone and the ionospheric parameters (Sect. 3.3). The power spectrum of JO^1D at Finokalia shows a peak at 32–45 min (99% significance level, not shown), while the UV irradiance at 305 nm spectrum at Heraklion, reveals two peaks at 28–45 min and 17–23 min. Similar oscillations were found in UV irradiance at distant stations from the peak eclipse, namely Ioannina and Nicosia (Kazantzidis et al., 2007).

Overall, a main oscillation with period in the range 30–40 min is found at a number of stations with independent measurements of total ozone column and its proxies JO^1D and UV irradiance, denoting an extended thermal stratospheric ozone disturbance.

3.3 Investigation of observed oscillations in the ionosphere

This solar eclipse took place under low geomagnetic and magnetospheric activities, providing clear advantage for the

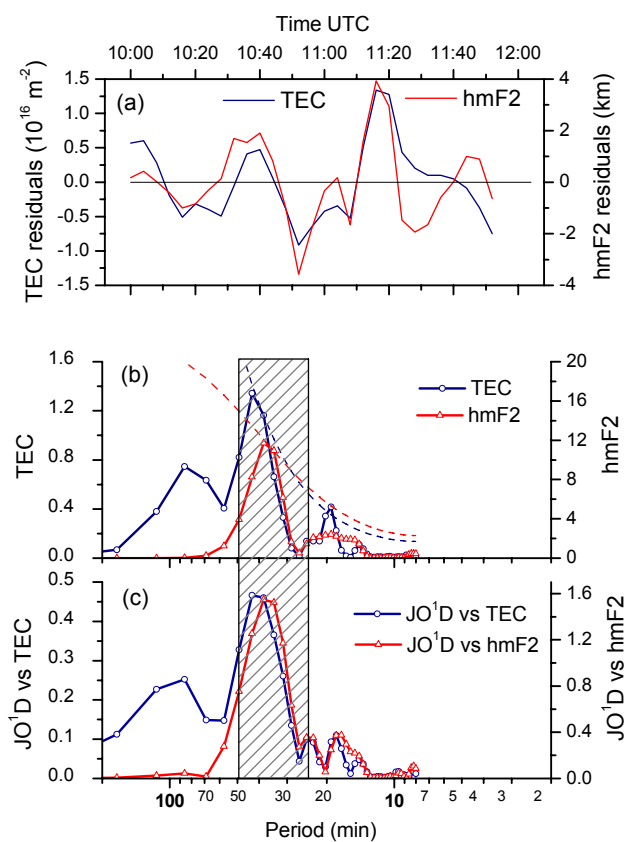


Fig. 4. (a) Ionosonde TEC and hmF2 residuals, (b) Spectrum Analysis applied on the residuals; Y-axes correspond to the spectral estimates of the parameters and similarly colored dash lines represent the 95% confidence limits of each spectrum, (c) Cross Spectrum Analysis between JO^1D at Kastelorizo and ionospheric parameters; Y-axes correspond to the amplitude of the cross-spectrum for the following pairs: JO^1D vs Ionosonde TEC (blue line + circles) and JO^1D vs hmF2 (red line + triangles).

identification of signals from solar eclipse induced GWs in the ionosphere (Fig. 3). In particular, the geomagnetic activity remained low, since the Dst index values ranged above -30 nT during the whole week (27–31 March 2006), and the same holds for the magnetospheric activity. Some moderate excursions in AE index (Mayaud, 1980) recorded in very early morning or very late evening hours the days prior to the eclipse day, cannot impose an effect in the ionosphere over Athens at the local time of the eclipse occurrence (Pröls, 1995).

The complete evolution of ionospheric parameters during the eclipse is thoroughly discussed by Gerasopoulos et al. (2007)¹. Here we focus on the oscillations observed in ITEC and hmF2 presented in Fig. 4a. Three main oscillations are found almost in coincidence with those of total ozone, however the lower time resolution of these measurements and the complexity of the exact GWs source (in the horizontal and the vertical), does not allow the estimation of

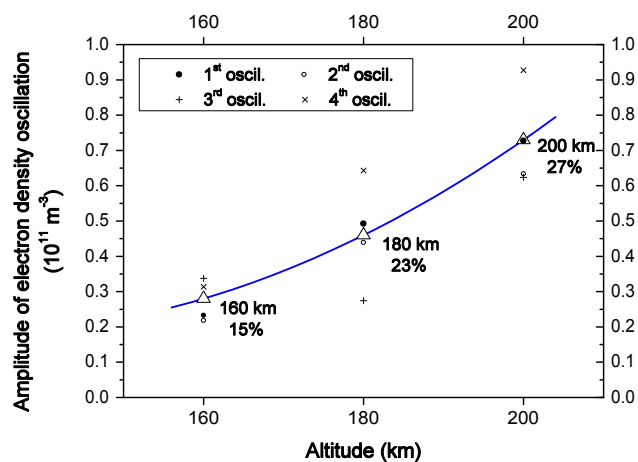


Fig. 5. Amplitudes of the electron density oscillations, calculated as the difference between successive minima and maxima at fixed altitudes (160, 180 and 200 km). Triangles represent the average amplitude of four oscillations at each altitude accompanied by the respective regression line (blue line), while each individual oscillation is shown with different symbol (see label). The percentage below the respective altitudes expresses the standardized to the average electron density (between first contact and maximum occultation) amplitude.

a reliable propagation velocity. The peak-to-peak amplitude of the ITEC residuals is 10–15% of the ITEC averaged over the eclipse period, while for hmF2 it is about 2%. The power spectra of ITEC and hmF2 are shown in Fig. 4b. Significant oscillations (95% confidence level) in the range 34–43 and 34–38 min are revealed for ITEC and hmF2, respectively. An additional peak is observed on the ITEC spectrum at 18–20 min. Power spectra were also calculated for all “quiet” geomagnetic conditions on days between 27 and 31 March 2006, and no significant fluctuations were found, denoting that the 29 March periodicities were indeed related to solar eclipse effects.

Cross-Spectrum Analysis has been applied between JO¹D at Kastelorizo and the ionospheric parameters. The amplitude of the cross spectra for the pairs JO¹D vs ITEC and JO¹D vs hmF2 is shown in Fig. 4c. A very distinct covariance between the frequency components corresponding to the ~40 min periodicity is found in both spectra. Significantly high square coherences were calculated, 0.93 and 0.96 for JO¹D vs ITEC and JO¹D vs hmF2, respectively, demonstrating that the ionospheric oscillations are probably driven by GWs initially formed in the stratosphere.

Variation of electron density amplitudes with height: Electron density variations at fixed ionospheric altitude zones of 20 km depth were additionally used to further speculate on the main source of GWs that have reached ionospheric heights. Consistent fluctuations with maximum amplitude in the altitude range of 200–220 km were observed during the first phase of the eclipse (Gerasopoulos et al., 2007¹). Such

oscillations were previously reported in the literature and are attributed to solar eclipse induced GWs during or after the solar eclipse (Altadill et al., 2001; Sauli et al., 2006). In this case, the fluctuations were considerably attenuated during the solar reappearance phase. The spectrum analysis (not shown) depicted a dominant periodicity of 15–18 min, coinciding with the higher frequency oscillations in total ozone and ITEC.

The change in the amplitude of these oscillations was used to speculate on the location of the GWs source responsible for the ionospheric signal. The amplitude of these oscillations was calculated as the difference between successive minima and maxima of the electron densities at the fixed altitude zones. These oscillations were clearly present in the altitude range 140–220 km. However, the 140–160 km zone was not included in our discussion here in order to keep the E-layer conditions clearly out of the analysis, since the electron concentration is very low and the detection of oscillations is ambiguous. Moreover, no such analysis was performed above 220 km since the main response of the ionosphere to the eclipse was gradually diminished from that height up (Gerasopoulos et al., 2007¹).

The amplitude of each of the four oscillations that can be identified with relative good precision, as well as the average amplitude per altitude is presented in Fig. 5. A tendency of increasing amplitude with altitude is clearly evidenced in both raw and standardized amplitudes, a result consistent with the effect of density decrease with height on the amplitude of a vertically propagating wave (Fritts and Luo, 1993).

In summary, taking into account that the observed oscillations i) are clearly present in the ionospheric height range 140–220 km, ii) are well attenuated above 220 km, iii) they have the same period at each height, iv) they have an important vertical propagation component, with increasing amplitude with height and v) are not of auroral origin, one could argue that they seem to originate from below the studied ionospheric heights. The above characteristics lie in favour of propagating waves attributed to the movement of the cooled spot produced by the moon’s shadow in the ozone layer.

3.4 Investigation of observed oscillations in the troposphere

The identification of GWs in the troposphere is a much more difficult task. Especially inside the boundary layer (BL), any fluctuations are subject to multiple rationales and increased uncertainty, related to meteorological or other local scale factors that may well mask the signal imposed by the propagation of GWs down to the surface.

We have used the temperature record at Kastelorizo to detect any fluctuations from GWs near the surface. The temperature drop of 2.3°C at Kastelorizo is shown in Fig. 6a (upper graph) accompanied by a polynomial fitting. Temperature residuals were calculated and are presented in Fig. 6a. (lower graph). The three oscillations seen in total ozone are also ob-

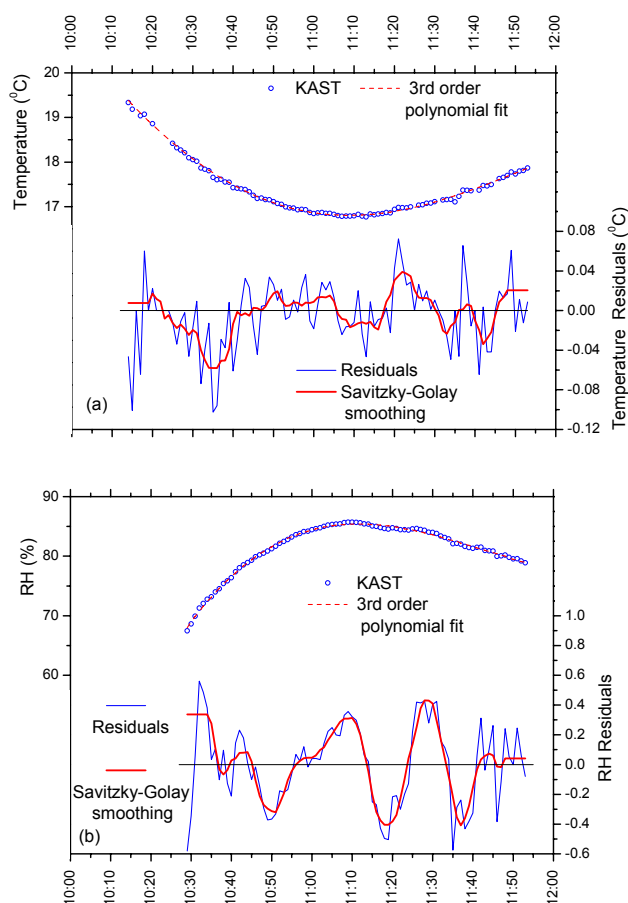


Fig. 6. (a) Upper panel: Temperature at Kastelorzio during the eclipse (open circles) and a third order polynomial fitted on the data (dash red line), (b) Upper panel: The same for RH at Kastelorzio. Lower panels in (a) and (b): Residuals of each parameter (blue line-dots) and a Savitzky-Golay (2nd order moving polynomial, 10 points) smoothing line (continuous red line). Time in UTC.

served in surface temperature with a 5–10 min lag, however the peak-to-peak amplitude of the residuals is rather small $\sim 0.1^\circ\text{C}$ or 0.6% of the temperature averaged over the eclipse period, which is same order of magnitude with the sensors accuracy ($\pm 0.1^\circ\text{C}$; Founda et al., 2007). This amplitude is one order of magnitude lower than that predicted by Eckermann et al. (2007).

To exclude relation of this signal to instrumental noise and overcome a possible lack of confidence in the temperature oscillations we have repeated the same analysis with relative humidity, RH (Fig. 6b). The residuals shown in Fig. 6b (lower graph) once more reveal the three dominant oscillations and the peak-to-peak amplitude is $\sim 1\%$ (absolute) or 1.2% of the RH averaged over the eclipse period. The accuracy of the RH sensor is $\pm 1\%$, so once more the amplitude of RH signal is comparable to the instrumental accuracy, however, the fact that the oscillations are very distinct and are observed in both temperature and RH enhances our confi-

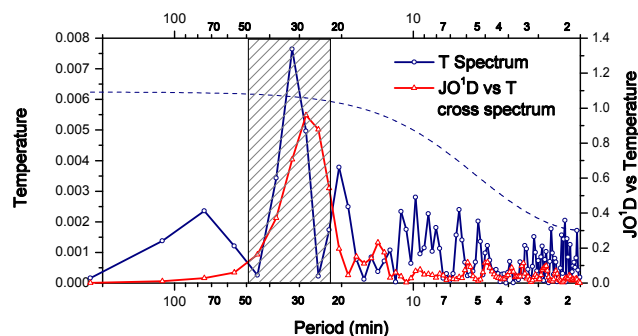


Fig. 7. Spectrum Analysis on temperature residuals and Cross Spectrum Analysis between JO^1D and temperature at Kastelorzio. Y-axes correspond to the spectral and cross spectral estimates, and the colored dash line represents the 95% confidence limits of the spectrum.

dence. It should be also noted here that averaging over 1 min intervals (measurements recorded every 20 s) has as a result to reduce the uncertainty both in temperature and RH measurements, and this is easily observed in Fig. 6a and b, where the noise is well separated from the oscillations.

The power spectrum of temperature at Kastelorzio is shown in Fig. 7 and a significant periodicity (95% confidence level) in the range 28–38 min is revealed. Similar oscillations, with modest wave amplitudes $\sim 0.1^\circ\text{C}$, were found in the spectra of the temperature time series at other sites namely Finokalia-Crete and Penteli-Athens (not shown). The cross spectrum between JO^1D and temperature (Fig. 7) shows a strong covariance around the 30 min periodicity with significantly high square coherence, 0.97.

Concluding, even though distinct oscillations are observed in surface temperature data at various sites, additionally re-confirmed by similar relative humidity periodicity, there is no clear evidence for GWs in the troposphere. Especially inside the boundary layer, manifold rationale could be provided for such transient perturbations in temperature and more parameters should be monitored and examined to draw safe conclusions.

4 Summary – conclusions

Definite observational evidence for a characteristic bow-wave response of the atmosphere to solar eclipse passages, despite model calculations at various altitudes, has been still ambiguous. In this paper, we have provided combined experimental evidence to support the initial hypothesis of Chimonas and Hines (1970) that the cooling of ozone layer in the stratosphere by the supersonic travel of the moon's shadow during an eclipse, constitutes a source of gravity waves propagating both upwards and downwards. To examine the above, concurrent measurements at three critical

layers in the atmosphere namely in the troposphere, in the stratosphere and in the ionosphere, were conducted.

Spectral Fourier Analysis revealed a dominant oscillation with period in the range 30–40 min in total ozone column over Kastelorizo (100% maximum occultation), but also at larger distances, with less sun coverage (Thessaloniki, 75% maximum occultation). It is expected that ozone in the upper stratosphere, where it maximizes and is in radiative equilibrium, would respond more rapidly to the transient cooling due to the eclipse of the sun. The eclipse induced ozone layer thermal disturbance was endorsed by JO¹D and UV irradiance (305 nm) measurements, both sensitive to columnar ozone variability, yet denoting a spatially extended propagation of the GWs with regard to the totality axis.

The 30–40 min oscillation was also evident in the spectra of the Ionosonde Total Electron Content (ITEC) and the peak electron density height in the ionosphere (hmF2). Cross-spectrum analysis between total ozone and the ionospheric parameters depicted high covariance in this range of periods. The initial argument that the source of the perturbation originates below the ionosphere was ratified by the fact that the amplitude of the electron density oscillation increased upwards from 160 to 220 km, which is expected for a vertically propagating wave inside a mean of decreasing with height density.

The identification of the GWs oscillation in the troposphere has been attempted with records of surface temperature and relative humidity. Distinct oscillations were observed in both parameters within the period range of our interest, at various sites. However, the amplitude of these oscillations has been modest and in the same order of magnitude with the sensor's accuracy. It should be kept in mind that the intensity of the waves downwards could be considerably suppressed by the fact that the propagation takes place in a denser mean, and that inside the BL any periodical signal in the parameters subjects to manifold rationale also controlled by meteorological or other local scale factors. The above do not allow us to draw safe conclusions on the influence of eclipse induced GWs in the troposphere, and should be taken under consideration for future experiments planning.

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